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First constraints on fuzzy dark matter from Lyman- α forest data and hydrodynamical simulations

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We present constraints on the masses of extremely light bosons dubbed fuzzy dark matter from Lyman- α forest data. Extremely light bosons with a De Broglie wavelength of ~ 1 kpc have been suggested as dark matter candidates that may resolve some of the current small scale problems of the cold dark matter model. For the first time we use hydrodynamical simulations to model the Lyman- α flux power spectrum in these models and compare with the observed flux power spectrum from two different data sets: the XQ-100 and HIRES/MIKE quasar spectra samples. After marginalization over nuisance and physical parameters and with conservative assumptions for the thermal history of the IGM that allow for jumps in the temperature of up to 5000 K, XQ-100 provides a lower limit of 7.1×10^{-22} eV, HIRES/MIKE returns a stronger limit of 14.3×10^{-22} eV, while the combination of both data sets results in a limit of 20×10^{-22} eV (2σ C.L.). The limits for the analysis of the combined data sets increases to 37.5×10^{-22} eV (2σ C.L.) when a smoother thermal history is assumed where the temperature of the IGM evolves as a power-law in redshift. Light boson masses in the range $1 - 10 \times 10^{-22}$ eV are ruled out at high significance by our analysis, casting strong doubts on suggestions of significant astrophysical implications of FDM, in particular for solving the “small scale crisis” of cold dark matter models.

Introduction. Recently, there has been a growing interest in so-called Fuzzy Dark Matter (FDM) models where the dark matter is made of ultra-light bosons. Cosmological and astrophysical consequences have been comprehensively reviewed in [1–3] highlighting the particle physics motivation [4, 5] for such models, as well as the importance of experimental searches [6]. A broad variety of astrophysical implications have been investigated in the literature: the halo mass function [7], the innermost structure of haloes [8], the dynamical properties of the smallest objects [9], the linear matter power spectrum [1], the development of non-linearities by using N-body simulations [10], the abundance of high redshift objects [11], the overall impact of FDM on galaxy formation and the reionization history of the Universe, the intergalactic medium [3, 12–14], pulsar timing and binary pulsars [15, 16], and the properties of our galactic disk [17]. The

general conclusion is that in order to have an appreciable astrophysical impact the mass of ultra-light bosons would have to lie in the range $1 - 10 \times 10^{-22}$ eV, and in this mass range it is indeed possible that some small scale “tensions” of cold dark matter with observations could be alleviated (e.g. [18] for a review of the small scale “crisis” of cold dark matter).

The intergalactic medium (IGM) [19, 20] plays a unique role in constraining the (small scale) matter power spectrum, since the low-density, high redshift IGM filaments are particularly sensitive to the small scale properties of dark matter. The main observable manifestation of the IGM, the Lyman- α forest, has provided important constraints on the linear matter power spectrum, especially when combined with cosmic microwave background data [21–27]. This includes, most notably, the tightest constraints on warm dark matter (WDM) mod-

els [28–33], upper limits on neutrino masses [26, 33, 34] as well as the recent remarkable discovery of Baryonic Acoustic Oscillations in the transmitted 3D flux [35, 36]. These results, especially those at small scales, are primarily due to the fact that the observed Lyman- α forest flux power spectrum is an integral of the 3D power spectrum and, and as the Lyman- α forest data probes high redshift, can probe small scale departures from the reference cold dark matter linear power spectrum more effectively than any other observable. At present, the tightest limits on the free streaming of WDM, expressed as the equivalent masses of thermal WDM relics, are in the range $m_{\text{WDM}} > 2.3 - 5$ keV at 2σ C.L.. The values at the lower end of this range are probably overly conservative and require the assumption of thermal histories that are likely unphysical [37]. The strongest constraints are obtained using SDSS-III/BOSS data [31, 33] and in particular the combined XQ-100 and HIRES/MIKE constraints [38] when restricting the thermal priors to physically plausible thermal histories. A robust, and slightly more conservative, lower bound of $3 - 3.3$ keV was suggested by both high resolution and low resolution quasar (QSO) spectra [29] as well as by [38] when allowing for a wider range of thermal histories of the IGM, including moderate size (5000 K) temperature jumps.

In FDM models, even though the typical Jeans scale is rather small (\sim kpc) the effect of FDM on the linear matter power spectrum is noticeable on scales larger than the smallest scales typically constrained by IGM data [1]. In the absence of fully numerical FDM simulations of the flux power spectrum, it has therefore become common practice to convert the limit on thermal relic WDM models – for which the flux power spectrum has been modelled in considerable detail – into FDM limits by comparing the linear matter power spectrum of WDM and FDM models and using the mass corresponding to $k_{1/2}$, the wavenumber at which the linear power spectrum departs (i.e. is suppressed) from the corresponding cold dark matter power spectrum by 50%. However, the accuracy of this rather crude mapping can only be checked by performing a full set of hydrodynamical simulations to model the effect of FDM on the properties of the IGM and the Lyman- α forest. Here, we will use such simulations to provide the first constraints on FDM models based on a full modelling of the Lyman- α flux power spectrum and comparison with two high-redshift data sets well suited to probing the small scale matter power spectrum. This will also allow us to check the accuracy of the $k_{1/2}$ mapping of thermal relic WDM constraints. Our analysis will be quite similar to the one presented in [29] and [38], where the flux power spectrum is modelled using a set of hydrodynamical simulations that vary astrophysical and cosmological parameters combined with a Monte Carlo Markov Chain analysis in the multi dimensional parameter space.

Data. The first sample we use is the set of 100 medium

resolution, high signal-to-noise QSO spectra of the XQ-100 survey [39], with emission redshifts $3.5 < z < 4.5$. A more detailed description of the data and the power spectrum measurements of the XQ-100 survey can be found in [38]. Here we repeat the most important properties of the data and the derived flux power spectrum. The spectral resolution of the X-shooter spectrograph is 30-50 km/s, depending on wavelength. The flux power spectrum used in the analysis has been calculated for a total of 114 (k, z) data points in the ranges $z = 3, 3.2, 3.4, 3.6, 3.8, 4, 4.2$ and 19 bins in k -space in the range 0.003-0.057 s/km. We further use the measurements of the flux power spectrum presented in [29], at redshift bins $z = 4.2, 4.6, 5.0, 5.4$ and in 10 k -bins in the range 0.001-0.08 s/km. In this second sample the spectral resolution of the QSO absorption spectra obtained with the MIKE and HIRES spectrographs are about 13.6 and 6.7 km/s, respectively. As in the analysis of [29], a conservative cut is imposed on the flux power spectrum obtained from the MIKE and HIRES data, and only the measurements with $k < 0.005 \text{ km}^{-1} \text{ s}$ are used to avoid possible systematic uncertainties on large scales due to the rather small number of QSO spectra in this sample.

Compared to XQ-100, the HIRES/MIKE sample has the advantage of probing smaller scales and higher redshift. There is a small redshift overlap between the two samples at $z = 4.2$. Since the thermal broadening (measured in km/s) of Lyman- α forest lines is roughly constant with redshift, the presence of a cutoff in the matter power spectrum due to free-streaming becomes more prominent in velocity space at high redshift due to the $H(z)/(1+z)$ scaling between the fixed comoving length scale set by the free-streaming length and the corresponding velocity scale. Moreover, the 1D power spectrum is more sensitive to the presence of a cutoff compared to the 3D power spectrum.

Simulations. Similarly as in [29] and [38], we model the flux power spectrum using a set of hydrodynamical simulations performed with the GADGET-3 code, a modified version of the publicly available Gadget-2 code [40]. A simplified star formation criterion is applied for which gas particles above an overdensity 1000, and below $T = 10^5$ K are converted into stars (e.g. [41]). The reference model simulation has a box length of 20 Mpc/ h with 2×768^3 and (cold) dark matter particles (gravitational softening of 1 comov. kpc/ h) in a flat Λ CDM Universe with cosmological parameters $\Omega_{\text{m}} = 0.301$, $\Omega_{\text{b}} = 0.0457$, $n_{\text{s}} = 0.961$, $H_0 = 70.2 \text{ km s}^{-1} \text{ Mpc}^{-1}$ and $\sigma_8 = 0.829$ in agreement with [42]. Three different WDM models with masses $m_{\text{WDM}} = 2, 3, 4$ keV have also been simulated in order to obtain WDM constraints.

We simulate 5 different FDM models using the transfer function provided by [1] with light axion masses m_{FDM} of 1, 4, 5.7, 15.7 and 30×10^{-22} eV. These values have roughly the same $k_{1/2}$ as models corresponding to thermal WDM relic masses of 1, 1.73, 2, 3, 3.87 keV, covering

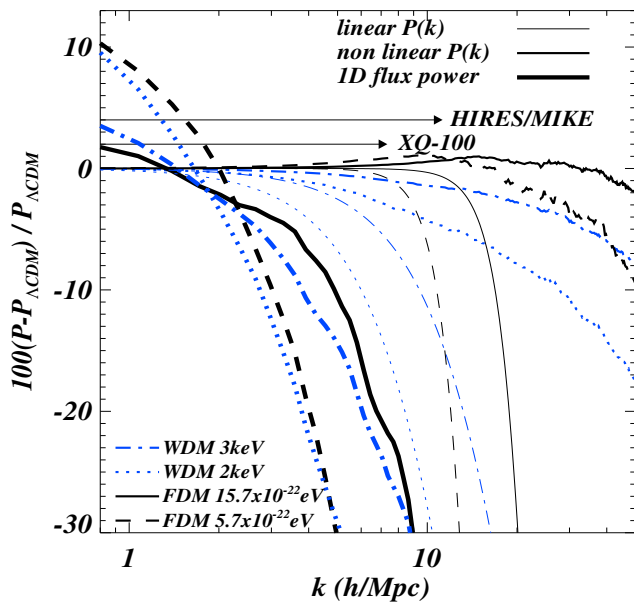


FIG. 1: Power spectrum relative to Λ CDM at $z = 5.4$ (in percent). Linear matter, non-linear matter and flux power spectra are represented by the thin, thick and very thick curves, respectively. Black (blue) curves are for FDM (WDM) models with $m_{\text{FDM}} = 5.7, 15.7 \times 10^{-22}$ eV ($m_{\text{WDM}} = 2, 3$ keV).

the range relevant for the “small scale crisis” of cold dark matter. These models were also simulated using the axionCAMB code [2] to obtain the linear transfer function, finding negligible impact on the simulated flux power. The corresponding Λ CDM model is also simulated along with a range of IGM thermal histories and cosmological parameters. In Fig. 1 we show the linear, non-linear and flux power spectrum at $z = 5.4$ for WDM and FDM models that have the same $k_{1/2}$: non-linearities erase some of the information contained in the linear power spectrum. Note that the 1D flux power is much more sensitive to the cutoff. The maximum wavenumbers at which the flux power spectrum is measured by Hires/MIKE and XQ-100 are represented by the horizontal arrows. It has been noted before (e.g. [43]) that for the analysis of the Lyman- α it is sufficient to use the appropriate transfer function without modelling the full quantum effects below the de-Broglie wavelength of the FDM particle.

We vary the thermal history by modifying the photo-heating rates in the simulations as in [44]. The low density IGM ($\Delta = 1 + \delta < 10$) is well described by a power-law temperature-density relation, $T = T_0 \Delta^{\gamma-1}$. We consider a range of values for the temperature at mean density T_0 and the slope of the $T - \rho$ relation, γ , based on the previous analysis of the Lyman- α forest and recent observations [45]. As in [38] these consist of a set of three different temperatures at mean density, $T_0(z = 3.6) = 7200, 11000, 14800$ K, which evolve with redshift, as well

as a set of three values of the slope of the $T - \rho$ relation: $\gamma(z = 3.6) = 1.0, 1.3, 1.5$. The reference thermal history assumes $(T_0(z = 3.6), \gamma(z = 3.6)) = (11000 \text{ K}, 1.5)$.

Following again [38] we use two parameters describing cosmology, σ_8 and $n_{\text{eff}} = d \ln P_m(k) / d \ln k$, evaluated at $k = 0.005 \text{ km}^{-1} \text{ s}$. Five different values are considered for both $\sigma_8 = 0.754, 0.804, 0.829, 0.854, 0.904$, and $n_{\text{eff}} = -2.3474, -2.3274, -2.3074, -2.2874, -2.2674$. The reference model has $(\sigma_8, n_{\text{eff}}, n_s) = (0.829, -2.3074, 0.961)$. We also vary the redshift of reionization z_{rei} which is chosen to be $z_{\text{rei}} = 9$ for the reference model as well as $z_{\text{rei}} = 7, 15$ for two additional models. The last parameter (f_{UV}) characterizes the effect of Ultraviolet (UV) background fluctuations. An extreme model dominated by QSOs has been chosen with a strong scale dependence at higher redshift and towards large scales. The mean flux is also varied a-posteriori through rescaling the effective optical depth, $\tau_{\text{eff}} = -\ln \bar{F}$. We use three different values $(0.8, 1, 1.2) \times \tau_{\text{obs,eff}}$, with the observed value of $\tau_{\text{obs,eff}}$ chosen to be those of the SDSS-III/BOSS measurements [46].

Method. Using the models of the flux obtained from the simulations we establish a grid of points for each redshift, in the parameter space of $(\bar{F}(z), T_0(z), \gamma(z), \sigma_8, z_{\text{rei}}, n_{\text{eff}}, f_{\text{UV}}, m_{\text{FDM}})$. We then perform a linear interpolation between the grid points in this multidimensional parameter space, to obtain predictions of flux power for the desired models. The interpolation is performed for $P_F(k, z)$, directly, rather than for ratios of flux power w.r.t. the corresponding Λ CDM simulation as was done in [29]. Parameter constraints are then obtained with a Monte Carlo Markov Chain code that explores the likelihood space until convergence is reached.

The redshift evolution of the IGM parameters T_0 and γ are modelled as power-laws for our reference analysis: $T_0(z) = T_0^A [(1+z)/(1+z_p)]^{T_0^S}$ and $\gamma(z) = \gamma^A [(1+z)/(1+z_p)]^{\gamma^S}$. The pivot redshift is different for each data set and roughly corresponds to the redshift at which most of the data lies ($z_p = 3.6, 4.5, 4.2$ for XQ-100, Hires/MIKE and the combined analysis, respectively).

Results. In Figure 2 we show the main result of this letter: the marginalized 1D likelihood for $1/m_{\text{FDM}}$. For our reference analysis, in which the temperature evolution is parameterized as a power-law at different pivot redshifts, XQ-100 returns an upper limit of 4.6×10^{-22} eV, Hires/MIKE gives 16.4×10^{-22} eV, while the combination of the two data sets results in a considerable improvement to 37.5×10^{-22} eV (2σ C.L.). These numbers become $2.7, 16.5, 32.2 \times 10^{-22}$ eV, for XQ-100, Hires/MIKE and both data sets, when using the following Planck priors on $n_{\text{eff}} = -2.307 \pm 0.01$ and $\sigma_8 = 0.829 \pm 0.01$ (1σ Gaussian priors). The improvement in the joint constraints is due to the fact that when combining the two data sets the thermal evolution of the IGM is

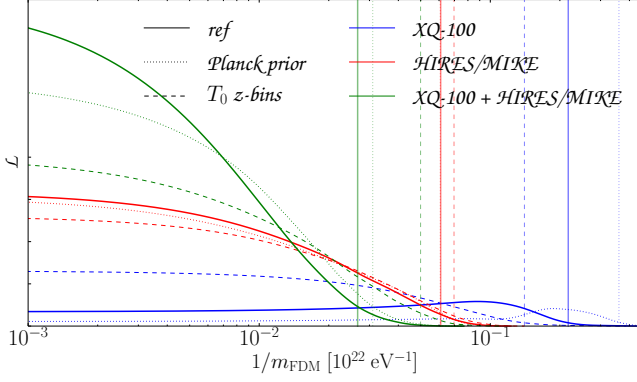


FIG. 2: 1D marginalized likelihood constraints for $1/m_{\text{FDM}}$. XQ-100, HIRES/MIKE and XQ-100+HIRES/MIKE are represented by the blue, red and green curves, respectively, with (continuous curves) and without (dotted curves) using Planck priors. Additionally results for models where we vary temperature independently in each redshift bin are plotted as dashed curves. 2σ upper limits are represented by vertical lines.

assumed to be one power-law for both $T_0(z)$ and $\gamma(z)$ for the full redshift range of the combined data sets. If we drop the assumption of a power-law evolution and we let the temperature vary independently in each redshift bin, with a maximum jump $\Delta T = 5000$ K in bins that are separated by $\Delta z = 0.2$, we obtain $7.1, 14.3, 20.0 \times 10^{-22}$ eV for XQ-100, HIRES/MIKE and both combined. We regard this result as the most conservative, since sudden jumps of temperature are not physically plausible in this redshift range (e.g. [47, 48]).

Increasing the covariance matrix by a multiplicative factor 1.3 in order to better represent a possible underestimation of the errors does not affect the results appreciably. In Tab. 1 we summarize the results including the $\chi^2/d.o.f.$ for the reference case, which appear to be very reasonable in all cases. For the combined analysis the other parameters lie within the following 2σ C.L. ranges at $z_p = 4.2$: $\sigma_8 = [0.83, 0.95]$, $n_{\text{eff}} = [-2.43, -2.31]$, $T^A(z_p) [10^4 \text{ K}] = [0.71, 1.06]$, $T^S(z_p) = [-3.37, -0.80]$, $\gamma^A(z_p) = [1.27, 1.69]$, $\gamma^S(z_p) = [-0.11, 1.82]$, $z_{\text{rei}} = [6.27, 13.62]$, $f_{\text{UV}} = [0.04, 0.94]$.

We have verified the constraints obtained by considering additional simulations with m_{FDM} of 30×10^{-22} eV.

In Figure 3 we illustrate the relation between WDM and FDM constraints. Due to the lack of the detailed modelling of the flux power spectrum for FDM models using hydrodynamical simulation so far, it has been usually assumed that one can map thermal relic WDM masses into FDM constraints by identifying the corresponding $k_{1/2}$ values for the linear power spectra, i.e. the wavenumber at which the power reaches 50% of the Λ CDM linear power spectrum. Figure 3 shows that this is not a good approximation. For XQ-100 only, i.e. for a

$m_{\text{FDM}} [10^{-22} \text{ eV}]$	XQ-100	HIRES/MIKE	Combined
ref.	4.5	16.4	37.5
Cov. $\times 1.3$	3.9	16.3	34.9
Planck priors	2.7	16.5	32.2
$T_0(z)$ bins	7.1	14.3	20.0
$\chi^2/d.o.f.$ (ref.)	134/124	33/40	187/173

TABLE I: Marginalized constraints at 95 % (lower limits). The pivot redshifts for different data sets are: $z_p = 3.6, 4.5, 4.2$ for XQ-100, HIRES/MIKE and combined for reference case, covariance matrix multiplied by 1.3, Planck priors and temperature in redshift bins.

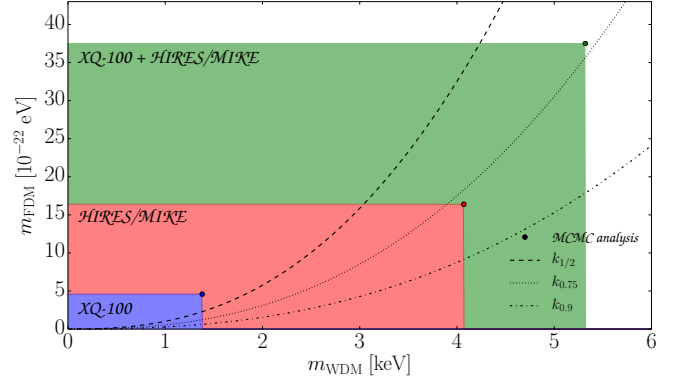


FIG. 3: Comparison of limits on m_{FDM} and m_{WDM} (thermal relic). Also shown is the relation obtained by matching $k_{1/2}$, $k_{0.75}$ and $k_{0.9}$ for the linear power spectra. Shaded areas show regions excluded by the analysis of XQ-100, HIRES/MIKE and both data sets in blue, red and green colors, respectively. Filled circles represent the 2σ C.L. lower limits for WDM and FDM.

WDM constraint of 1.34 keV, the corresponding FDM limit is slightly above the $k_{1/2}$ curve, which remains a reasonable approximation. The HIRES/MIKE data only, however, gives a WDM lower limit of 4.7 keV that translates into a FDM limit which is much weaker than the one that would be obtained using the $k_{1/2}$ mapping. The same holds for the combined analysis. For these higher thermal relic WDM masses the WDM constraints are better mapped into FDM constraints by using $k_{0.75}$ rather than $k_{1/2}$, where $k_{0.75}$ is defined as the wavenumber at which the power reaches 75% of the Λ CDM linear power spectrum. The reason for this is that as the free-streaming cut-off moves to smaller scales for larger particle masses, the scales affected by the free-streaming cut-off become more non-linear.

We have also checked that similar conclusions are reached when considering the 1D linear power spectrum, which is the quantity really probed by the flux power spectrum (rather than the 3D power spectrum). The different data sets are obviously constraining linear power

at different (k, z) values. Non-linearities will develop differently in WDM and FDM models and this makes the mapping between the two scenarios not straightforward. Full non-linear hydrodynamic simulations and detailed modelling of the flux power spectrum is required for accurate constraints on FDM models. WDM models retain more small scale power at $k > k_{0.75}$ compared to the corresponding FDM model that has instead a more prominent knee at $k < k_{0.75}$. These differences partially compensate in terms of non linear 1D flux power in a non-trivial way that depends on how non-linear the matter power spectrum is at the free-streaming scale. Similar to the shape of the WDM cutoff, the FDM cutoff in the flux power spectrum appears to be rather distinctive with no significant degeneracies with the other parameters in the analysis.

Conclusions. We have presented constraints on FDM models based on detailed modelling of the Lyman- α forest 1D flux power spectrum and high resolution data at intermediate and high redshifts with hydrodynamical simulations. These are the first constraints that incorporate the effect of the relevant IGM physics, including thermal and pressure smoothing on the non-linear evolution of the flux power spectrum on the relevant scales. Our final, conservative lower limit from a joint analysis of the intermediate and high-resolution Lyman- α forest data, $m_{\text{FDM}} > 20 \times 10^{-22}$ eV (2σ C.L.), was obtained with conservative assumptions for the thermal history of the IGM that allow for (unphysical) sudden jumps of the IGM temperature up to 5000K. This lower limit for the mass of ultra-light bosons strengthens by about a further factor two if we assume a smoother thermal history of the IGM. Our analysis appears to close the window of FDM models with significant astrophysical implications, in particular for alleviating the tension between observations and theoretical predictions of cold dark matter models on small scales.

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